

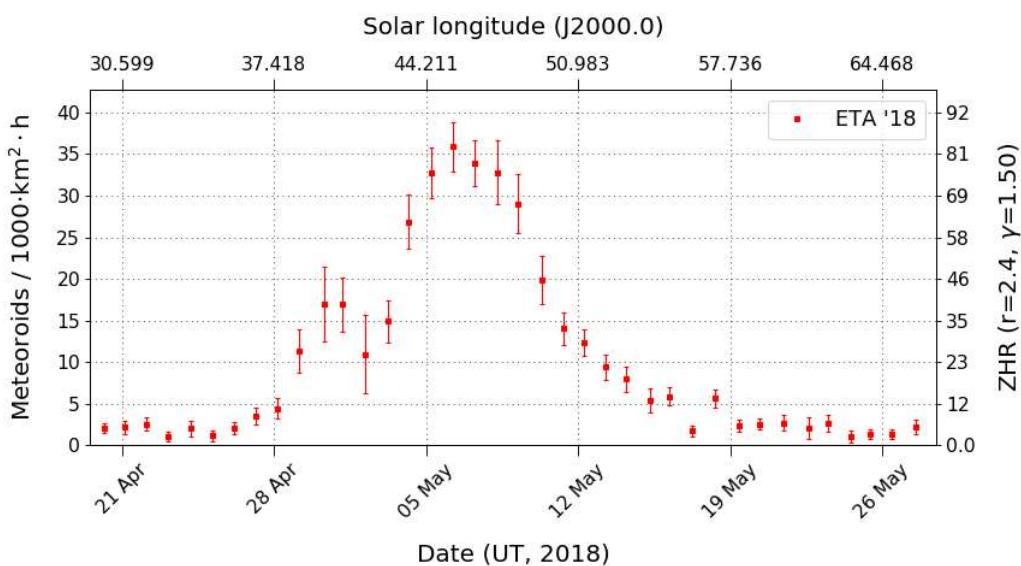
## Results of the IMO Video Meteor Network – May 2018

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2019/05/25

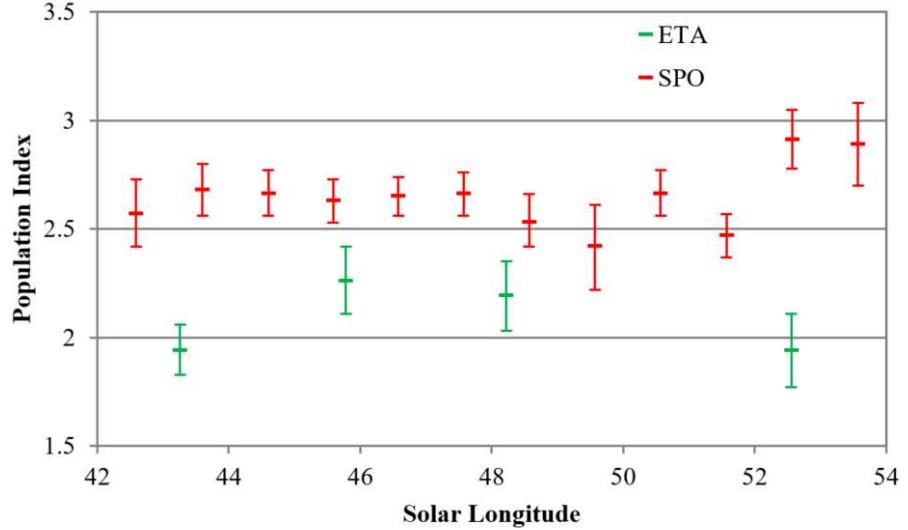
There is no other month where the climatic conditions are as stable as in May. At least we see that the observation results varied only little in the past four years. The effective observing time fluctuated between 7,000 and 7,800 hours, in which we recorded between 16,500 and 18,300 meteors. Those 7,500 observing hours which 40 observers collected with 78 video cameras in May 2018, match perfectly to the average. Only the yield was a little lower with 15,000 meteors. A quick look at the observation table teaches us that the first half of May was slightly better than the second. 80% of the cameras managed to observe in twenty or more observing nights, which is a top-class result. All observers but those from Slovenia enjoyed great observing conditions.

The eta Aquariids are the highlight of May, and we have reported about this shower several times. Whereas we notice only little from this shower in Central Europe, it is *the* shower of the year in the southern hemisphere. Figure 1 shows the activity profile of the eta Aquariids in 2018 covering the whole activity period. Around April 27 the activity starts to raise, and at the border between April and May it has already risen to 10 meteoroids per 1,000 km<sup>2</sup> and hour. Between May 5 and 9 we see a distinct plateau of high activity with over 30 meteoroids per 1,000 km<sup>2</sup> and hour. Thereafter the activity is quickly declining and reaches the base level near May 18. There are remarkable fluctuations early May, which can be attributed to an insufficient dataset between May 1 and 3, though.



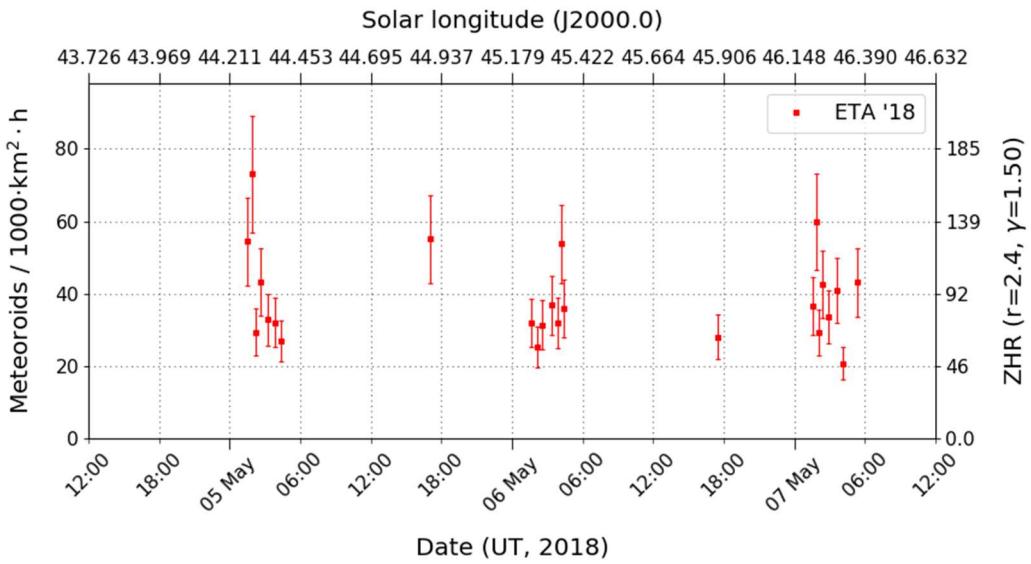
**Figure 1:** Flux density profile of the eta Aquariids in May 2018, derived from video data of the IMO Network.

The population index of the eta Aquariids can only be determined over longer time intervals, because even in the maximum nights our cameras record too few meteors. Near the peak it has a value of about  $r=2.2$ , whereas at the same time the sporadic meteors show a population index of  $r=2.6$  (figure 2). Last year we had obtained similar  $r$ -values of 2.0 and 2.6, respectively.



**Figure 2:** Population index profile of the eta Aquariids and sporadic meteors in May 2018.

Because of their radiant position, the eta Aquariids are of particular interest. Those few shower members which can be observed from central Europe, occur always at low radiant altitudes. Effects like the zenith exponent, which have a strong impact on the flux density at low radiant altitudes, can be analysed particularly well with this shower. Figure 3 shows three peak nights in detail. We do not see any systematic variation, which implies that the zenith exponent of 1.5 is of the right order.

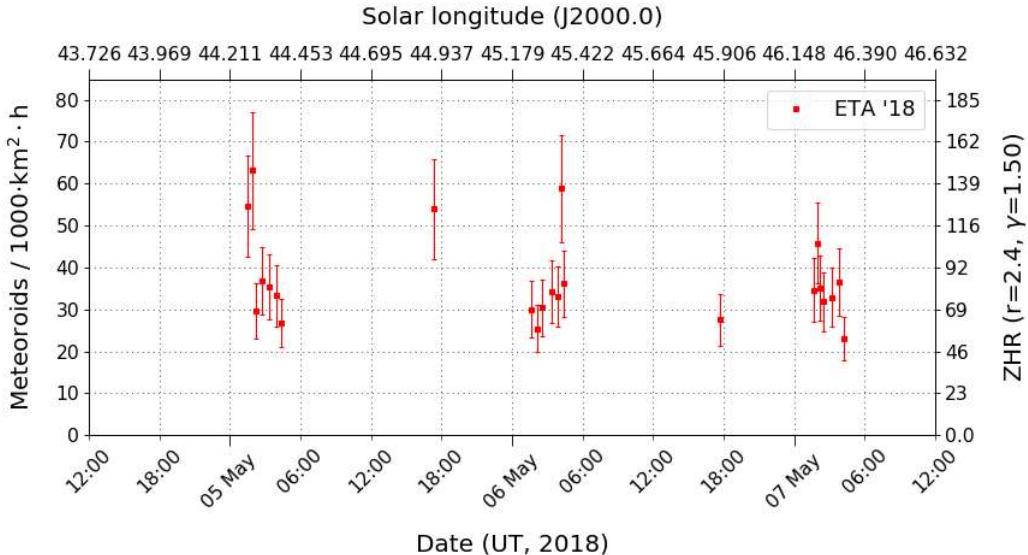


**Figure 3:** High resolution flux density profile of the eta Aquariids on May 4-7, 2018.

On the other hand, most observations are hampered by twilight, which leaves some uncertainty if that has an impact on the analysis. We could exclude that with an additional selection criterion in MeteorFlux. We could define a minimum stellar limiting magnitude, for example, to filter out observation at dawn. In addition, we could configure a maximum radiant altitude beside the already existing minimum altitude, to analyse other showers like the Quadrantids, Perseids or Geminids under similar radiant altitudes.

Unfortunately, both programmers that developed MeteorFlux and migrated in on a new AWS instance, have pulled themselves out of the project a long time ago. Many of my change requests have been put on hold since 2013, and since I have no knowledge of JavaScript, Python and PostgreSQL I never dared to approach the code.

During the ETA analysis of 2018, however, I decided to have a look at the source code, anyway, and to implement these new filters by copy&paste from existing code fragments. The result was quite encouraging, since after two evening I had understood the rough structure of the code and successfully implemented the additional filters. Figure 4 shows the same flux density profile as before, but only including observing intervals with a stellar limiting magnitude of 2 mag or better. The flux density profile changes only a little, so twilight does not seem to have a significant impact on the activity profile.



**Figure 4:** High resolution flux density profile of the eta Aquariids on May 4-7, 2018, using only intervals with a stellar limiting magnitude better than 2.0 mag.

Spurred on this success, I dared to implement the next selection criterion, which I had been waiting for many years, right away: the option to select individual cameras. This functionality was implemented in one evening, and later I provided the option to select the camera set by the observer, country and continent.

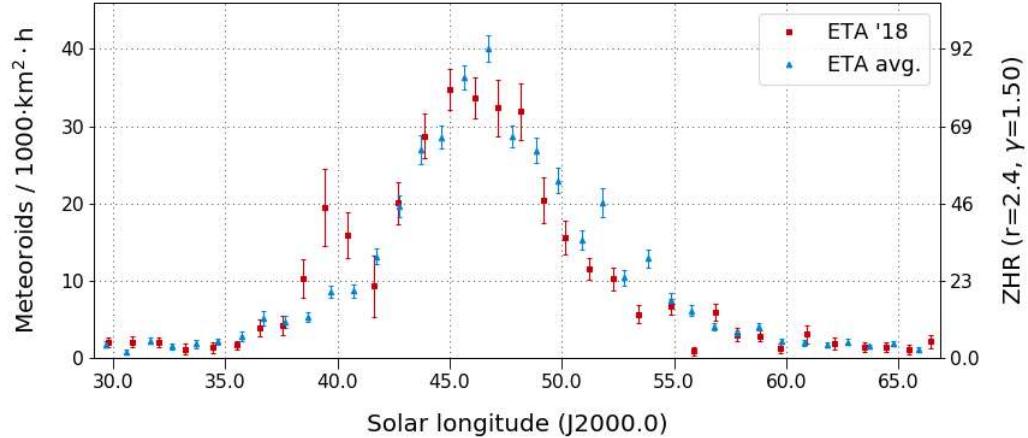
Euphorically I addressed a third aspect thereafter. Often you do not only want to create a single activity graph, but you want to compare two flux density profiles with one another. The aforementioned fluctuations of the eta Aquariids early May 2018 are a nice example. We want to compare the profiles with the average of the previous years to see if this is a recurring structure. Up to now I had to use a trick by generating both graphs independently of each other, and merge them together with Photoshop. Now you can generate two profiles with MeteorFlux in a single graph. In fact, you can not only vary the observing year, but also the other parameters used to generate the reference profile. This allows for a range of new analysis options, which we want to use intensively in the future, e.g.:

- You can compare the profile of a meteor shower from one year with the average of other years (and adapt the binning of the reference profile to lower or higher temporal resolution depending on the meteor number).
- The mean activity profile of years with new moon and full moon can be compared to search for systematic deviations.
- You can compare the activity of two meteor showers in the same time interval, e.g. from a meteor shower and sporadic meteors, or from the northern and southern Taurids. For better visibility, the reference shower can be offset and scaled linearly.
- If you select different binning parameters, you can compare a low resolution and high-resolution activity profile.
- Every observer can compare the results of his camera(s) with the outcome of other cameras. You can also compare the result of cameras from different countries.

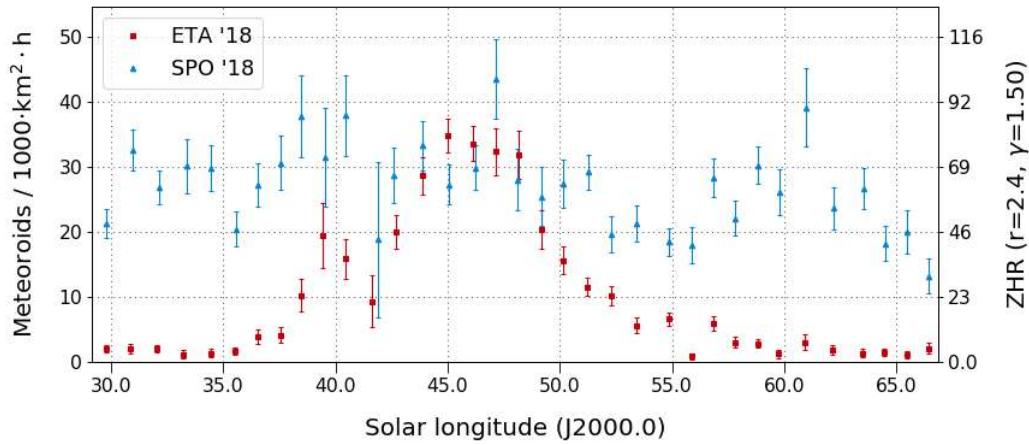
- You can compare data with good and poor limiting magnitude, or observations with low and high radiant altitude.
- The impact of different zenith exponents can be studied directly in a single graph.

Also, these functional extensions were implemented in three evenings. The following figures 5-7 give a few examples for the new functions, which can now be used by everyone at [meteorflux.org](http://meteorflux.org).

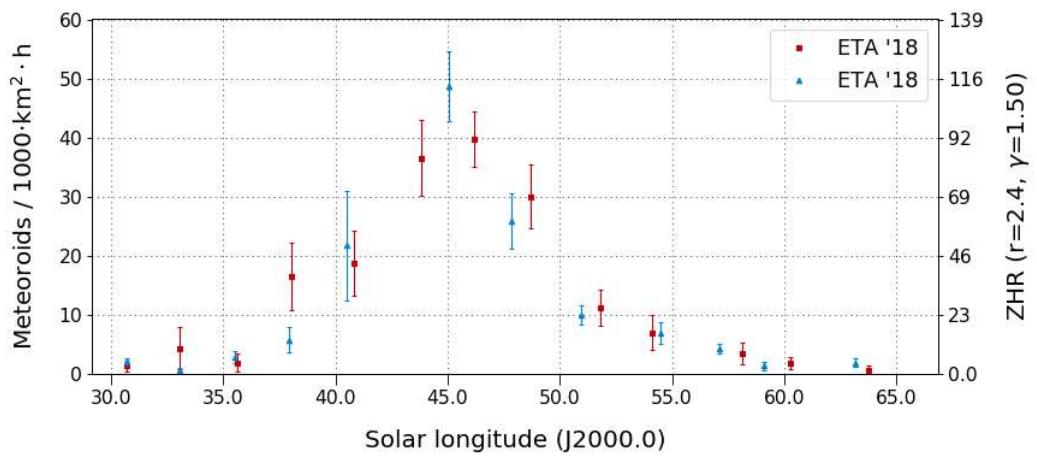
So far, all functions have been implemented primarily by copy&paste, but now I'm sufficiently optimistic to also try step by step extensions which require some new code.



**Figure 5:** Flux density profile of the eta Aquariids in 2018 (red) and in the average of 2014-2017 (blue), derived from video data of the IMO Network.



**Figure 6:** Flux density profile of the eta Aquariids (red) and sporadic meteors (blue) in May 2018.



**Figure 7:** Flux density profile of the eta Aquariids in May 2018, recorded by video cameras in Germany (red) and Italy (blue).





May	16	17	18	19	20	21	22	23	24	25	26	27	28	29	30	31
ARLRA	4.1	4.6	-	4.7	4.6	4.4	4.4	4.5	4.3	1.4	2.6	2.9	4.0	3.9	4.1	3.7
BERER	-	-	-	-	-	-	-	-	-	-	-	-	-	-	-	-
BIATO	7.6	2.6	7.5	7.5	5.6	-	1.6	7.1	-	7.1	7.1	3.7	-	0.2	-	3.6
BOMMA	7.5	1.1	7.4	7.4	5.1	-	1.9	6.3	4.6	7.0	7.0	3.5	0.2	0.9	0.3	6.7
BREMA	5.0	-	1.9	6.2	6.3	3.6	-	6.0	-	3.3	5.8	-	5.5	-	3.7	4.0
BRIBE	0.5	1.6	5.9	5.8	5.7	0.2	-	3.9	0.7	5.4	5.3	1.1	5.1	0.2	2.2	3.1
CARMA	1.0	4.6	1.2	5.4	4.8	3.6	-	-	2.3	5.0	5.3	1.5	1.8	2.4	-	0.9
CASFL	-	4.7	6.3	6.9	3.0	-	-	-	6.7	1.3	-	-	-	2.5	0.8	5.3
CINFR	-	4.8	5.6	7.0	3.2	-	-	0.6	6.6	2.4	2.8	-	-	2.1	1.1	5.6
CRIST	4.1	0.3	4.0	3.7	2.7	-	-	7.1	6.4	7.1	7.1	4.4	-	-	-	-
GONRU	1.3	5.8	-	3.6	6.9	-	2.1	5.0	6.7	6.7	5.0	0.6	-	1.7	1.1	1.1
ELTMA	2.8	7.0	-	5.1	6.9	-	0.8	5.1	6.0	4.8	3.7	0.6	-	0.7	1.1	0.7
FORKE	1.0	5.7	7.0	4.5	6.4	-	2.1	5.6	1.0	6.7	6.4	1.1	-	0.6	2.1	-
GOVMI	3.3	7.0	7.0	5.1	6.9	-	2.9	6.8	6.7	6.7	6.1	1.4	0.2	2.1	2.1	1.1
HERCA	5.3	7.1	6.8	7.2	3.8	0.8	-	-	3.3	-	-	-	-	0.3	0.7	2.9
HINWO	-	-	-	1.8	5.9	5.8	-	-	-	1.8	2.7	0.2	1.7	-	4.8	-
IGAAN	-	-	7.3	-	-	-	-	-	-	-	-	-	-	-	-	-
JONKA	7.8	7.8	7.8	6.1	4.6	6.5	7.6	1.9	1.1	4.2	1.0	-	3.4	1.3	-	4.4
KACJA	7.9	7.7	7.8	6.2	-	6.6	7.7	0.8	1.5	2.9	0.9	-	2.4	-	-	4.1
KACJA	7.8	7.4	7.0	5.9	4.1	6.0	7.2	0.5	-	5.0	-	3.2	5.8	-	-	4.0
MARRU	8.0	7.9	7.7	5.0	3.9	5.5	4.9	0.2	0.2	-	-	-	-	1.5	-	3.7
MOLSI	7.7	2.5	7.2	5.7	4.2	6.3	7.4	-	-	4.5	-	3.5	4.3	-	-	3.4
MORJO	4.6	1.5	3.2	5.3	2.3	6.4	1.3	-	1.8	6.3	2.3	3.8	0.4	-	4.9	1.2
MOSFA	-	0.8	0.2	1.8	0.2	3.6	3.8	3.9	-	0.5	6.1	3.8	-	0.3	-	3.0
NAGHE	0.3	-	1.4	5.0	4.2	3.7	0.8	4.5	0.4	4.9	4.1	4.9	5.8	1.9	2.9	1.5
OCHPA	1.6	-	0.6	-	6.1	1.2	0.9	2.8	0.2	0.8	2.1	2.1	3.5	2.9	1.8	2.1
OTTMI	-	4.8	-	-	-	2.8	-	-	3.4	2.3	-	-	-	2.8	-	4.6
PERZS	7.3	-	-	-	-	4.9	7.4	1.2	0.2	7.3	6.7	6.9	-	5.0	7.0	7.1
ROTEC	6.5	-	2.7	-	6.6	6.8	6.0	-	3.0	5.9	6.6	4.0	-	-	2.5	-
SARAN	-	5.2	-	5.3	5.3	5.2	5.1	4.8	4.8	4.3	-	-	0.7	4.3	4.4	0.8
SCALE	8.1	8.1	3.9	8.1	4.1	-	-	1.3	3.0	5.1	2.5	0.9	-	-	2.8	5.6
SCHHA	8.0	-	-	7.9	5.1	-	0.6	0.5	-	-	6.1	1.6	-	-	3.6	5.2
SLAST	7.8	-	-	7.9	5.4	-	1.2	0.9	-	-	6.6	2.5	-	-	3.3	5.8
STOEN	-	-	-	7.4	-	-	0.7	-	-	4.0	-	-	-	4.4	2.8	-
STRJO	7.6	5.9	-	7.8	2.4	-	-	0.8	1.2	4.8	2.0	-	-	-	1.2	-
TEPIS	3.9	5.0	5.8	6.3	3.1	-	2.1	-	4.3	2.3	1.4	-	-	2.9	3.1	6.3
WEGWA	3.4	2.0	5.5	5.7	5.6	2.5	-	4.1	3.3	0.3	0.3	2.4	4.6	4.0	3.0	-
YRJIL	0.6	4.2	2.6	3.6	1.4	3.6	-	-	3.9	2.8	-	-	0.3	0.3	-	0.6
ZAKJU	-	5.6	-	-	-	4.4	-	-	2.3	3.0	-	-	-	-	4.1	3.3
Sum	204.9	200.9	253.1	332.0	303.9	257.9	153.8	176.5	172.5	255.4	219.1	166.7	188.0	109.0	192.6	226.3



